

Interactive comment on “Universal multifractal Martian topography” by F. Landais et al.

Anonymous Referee #1

This manuscript is well written and provides new results. I have several rather minor comments:

- P. 1012, line 24: the parameter α is estimated using a second derivative. Furthermore it is a local estimate, around the moment $q=1$. The authors could estimate it by a best fit among all values admissible for α , chosen over a given range (moments q in $(0,2)$).

answer : We thank the referee for his positive evaluation. The short answer is yes, there are other ways of estimating α . We do however need to be cautious since due to multifractal “phase transitions”, the universal form breaks down for q larger than a critical value beyond which, $K(q)$ becomes linear. This is why we only considered $0 < q < 2$.

- It seems that the analysis of this 2D topography field is done along 1D cuts. Why? Could the authors perform analyses spatially in 2D? This is possible using structure functions. Is it possible using Haar wavelets?

answer : The study of anisotropy is certainly a good idea. Haar wavelets may be generalised in 2D but given the fact that MOLA data are not a regular grid of points, a global 2D analysis would require a considerable amount of time. Furthermore, the study of anisotropy is different from the multifractal property that we would like to focus on.

- P. 1017 line 18. The value $H=0.75$ for small scales could be related to the fracture value which was found to be close to 0.80 (e.g. Schmittbuhl et al, JGR, 1995). Some discussions on this point may be added in the manuscript.

answers : We added this discussion in the manuscript

- Typo P. 1017, line 17: "with": there are here some missing words

answers : We simply removed the word “with” placed here by mistake

- Table 1: negative values for α and $C1$: this is not possible and should be removed. The text indicates that the experimental curve is linear, hence these values are not relevant and only the H value should be given.

answers : we have replaced the irrelevant value of α by NaN

- Figs. 2, 5, 7: C302 please put the units for horizontal axis, and instead of \log_{10} linear scale, choose real values and a logarithmic scale (which is much more legible).

answer : ok

Interactive comment on “Universal multifractal Martian topography” by F. Landais et al.

Anonymous Referee #2

The paper is technically sound and the results are convincing. My main objection is instrumental. If MOLA samples are separated 300 m and the accuracy is 1.8 m (some authors say worse), appreciating the right scaling for differences at the scale of few kilometres sounds as challenging to me. How can we be sure that the effect below 10 Km is not merely an artifact (especially when the changes in slope do not seem so big in Figure 2)? One possible way to analyse this could be to rebin the series, for instance getting a series with half samples and each sample being the average of two adjacent original samples; this way the error is reduced by a factor square root of two, although the resolution is decreased by a factor of 2. If the crossover is not modified after this change this would imply that it is an actual geophysical limit; however, if it is decreased this would mean that the effect is an artifact. As this is an important point of the paper, I would like to see this question clarified prior to my acceptance of the paper, which is otherwise of great quality

Answer :

We thank the referee for his positive evaluation. As mentioned in your comment, using simple differences to appreciate the scalings at small scales may be misleading due to the limit in instrumental accuracy. For this reason precisely, we have not used simple differences to define fluctuations. As specified in part 3, we choosed the Haar fluctuations that have the advantage of reducing the uncertainty at small scales. Indeed, due to the definition of this wavelet, the smallest scale available is 600m (twice the MOLA resolution as you suggested) and the corresponding fluctuation is obtained from the data of 4 along-track points by taking the difference between the mean elevation of the first two points and the mean elevation of the last two points. This being equivalent to the procedure you suggested, we can not use your idea to test the accuracy of our results. However, It is clear to us that a bias at small scale is possible and should be carefully examined. To make sure that the accuracy of MOLA is not responsible for the transition around 10 km, we performed numerous fractal simulations artificially polluted by a random white noise to imitate the uncertainty of MOLA’s measurement. Our results indicated that this kind of error is very unlikely to produce the observed transition.



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This discussion paper is/has been under review for the journal Nonlinear Processes in Geophysics (NPG). Please refer to the corresponding final paper in NPG if available.

Universal multifractal Martian topography

F. Landais^{1,2}, F. Schmidt^{1,2}, and S. Lovejoy³

¹Univ Paris-Sud, Laboratoire GEOPS, UMR8148, Orsay, 91405

²CNRS, Orsay, 91405

³Physics department, McGill University, 3600 University st., Montreal, Que. H3A 2T8, Canada

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Correspondence to: F. Landais (francois.landais@u-psud.fr)

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Abstract

In the present study, we investigate the scaling properties of the topography of Mars. Planetary topographic fields are well known to roughly exhibit (mono)fractal behavior. Indeed, the fractal formalism reproduces much of the variability observed in topography. Still, a single fractal dimension is not enough to explain the huge variability and intermittency. Previous studies have claimed that fractal dimensions might be different from one region to another, excluding a general description at the planetary scale. In this article, we are analyzing the Martian topographic data with a multifractal formalism to study the scaling intermittency. In the multifractal paradigm, the apparent local variation of the fractal dimension is interpreted as a statistical property of multifractal fields. We analyze the topography measured with the laser altimeter MOLA at 300 m horizontal resolution, 1 m vertical resolution. We adapted the Haar fluctuation method to the irregularly sampled signal. The results suggest a multifractal behavior from planetary scale down to 10 km. From 10 km to 300 m, the topography seems to be simple monofractal. This transition indicates a significant change in the geological processes governing the Red Planet's surface.

1 Introduction

The acquisition of altimetric data from Mars Orbiter Laser altimeter (MOLA) has motivated numerous analyses of the Martian topography, each one aiming to properly characterize the surface roughness. A possible approach is to assume that topography can be statistically described with quantitative parameters able to characterize the geological units. Many statistical indicators have been proposed and widely explored in order to study the surface of Mars: RMS height, RMS slope, median slope (Kreslavsky and Head, 2000), autocorrelation length (Aharonson et al., 2001) (see Shepard et al., 2001 for a review of all those indicators). Useful information has been obtained by the use of those indicators but they have the disadvantage of being defined

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at a given scale. By construction, they do not directly take into account the scale symmetry that often occurs in the case of natural surfaces (Mandelbrot, 1967; Turcotte, 1997; Gilbert, 1989; Pommerol et al., 2012).

On Mars, different authors (Nikora and Goring, 2004) have explored the scaling properties of topography by the use of scale invariant parameters. Malamud and Turcotte (2001) has performed a wavelet analysis to study the polar topography of Mars. Other studies performed more classical wavelet analyses in the spatial domain: following the procedure proposed by Shepard et al. (2001); Orosei et al. (2003) used a quantity called the RMS deviation computed at different scales and exhibiting a power-law dependance on scales with a scaling exponent H (for Gaussian processes, the Hurst exponent). The value of H , evaluated at different locations was found to be different from a region to the other, revealing a strong intermittency in the statistical variability of the surface of Mars (Aharonson et al., 2001; Orosei et al., 2003). The observed intermittency, also found in the case of Earth (Baldassarri et al., 2008) and less pronounced in the case of the moon, (Rosenburg et al., 2011, 2015), might lead one to reject the idea of a global description of any topographic field at the planetary scale. However, modern developments in the theory of scale invariant processes might be able to give a full account to the observed variability and intermittency. As proposed by Lavalley et al. (1993), it is possible to extend the fractal interpretation of topography to a multifractal object requiring an infinite number of fractal dimensions (one for each altitude level). Multifractal simulations performed by Gagnon et al. (2006) revealed that synthetic multifractal fields tend to reproduce very convincingly the variability and intermittency of natural surfaces. A multifractal study has also been performed on artificial topographic DEMs revealing that multifractal parameters may be able to discriminate between different stages in topographic evolution (Vidal Vazquez et al., 2008). Moreover a particular class of multifractal called the universal multifractal has been proposed by Schertzer and Lovejoy (1987) (see Sect. 2.2). Stable and attractive, universal multifractals are good candidates to model the variability of topographic fields and have the advantage of being simply

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characterized by three parameters (see Sect. 2.2). In the case of Earth, this formalism has proven to be relevant (Gagnon et al., 2006) providing satisfying fit of the structure function computed in different area of the globe and allowing to quantify multifractal behavior at the planetary scale through the measurement of three parameters: H (the conservation exponent or fluctuation exponent), α (the degree of multifractality) and C_1 (the codimension of the mean). In this study, we aim to measure the global value of H , α and C_1 in the case of Mars. After a careful choice in the definition of fluctuations computed from all the available data in the MOLA database, statistical moments are computed on a large range of scales in order to estimate the universal multifractal parameters.

2 Methodology

This section first describes the Martian topography dataset used in this study and the MOLA instrument. The second part contains some elements of the theory of universal multifractality. The third part contains the description of the Haar fluctuation tool we adapted in case of the irregularly sampled Martian topography.

2.1 MOLA laser altimeter

MOLA (Mars Orbiter Laser altimeter Smith et al., 2001) is a laser altimeter on board of Mars Global Surveyor (NASA) spacecraft following a quasi-polar and circular orbit. MOLA has recorded a huge amount of topographic data providing a well-detailed global mapping of the planet. With 20 measurement every two seconds, a vertical accuracy of 1 m, a surface spot size of 168 m and approximatively 300 m between two consecutive measurements, the MOLA database stored in PDS (Planetary Data System, <http://pds-geosciences.wustl.edu>) constitutes an ideal dataset to study the global properties of Mars' topography down to the kilometer scale. Topography is calculated using the gravity field from Goddard Mars Model 3 (mgm1025, Lemoine et al., 2001). The global

areoid error through degree 60×60 for the mgm1025 model is 1.8m. The total number of individual topographic measurements used for the purpose of this study is close to 600 millions. For methodological reasons, we have decided to limit the scope of our analysis to the along-track series, meaning that we exclude all kind of across-track fluctuations (measurements extracted from different orbits). This restriction might introduce a bias, consequence of the single preferred direction (close to North/South) of MOLA's orbit. Hence, anisotropic statistical properties if they occur will not be detected by our method. However the assumption that topography is isotropic at the planetary scale seems reasonable.

2.2 Universal multifractality

The topography is a two-dimensional field providing for each pair of latitude/longitude a elevation value. In order to investigate the scaling properties of such field, we must study the distribution of slopes at different scales or equivalently, the fluctuations of elevation from one point to the other. The simplest way to define fluctuations is to compute the first difference for each couple of elevation data. The formalism used on this study is largely based on Lovejoy and Schertzer (2013).

Once fluctuations are defined, a common way to explore scaling properties of any geophysical field is to compute statistical moments of several orders and at different scales. If the field is scaling and if the fluctuations are defined properly, statistical moments will appear as straight lines on log-log plots, meaning they follow a power-law dependency on scales. This property is expressed by Eq. (1) where Δh stands for fluctuation of elevation and Δx stands for scale. In the case of simple fractality (i.e., *monofractality*), the q dependancy of $\zeta(q)$ is purely linear: $\zeta(q) = qH$ where H is called the conservation exponent used in honour of Edwin Hurst (typically a fractional brownian motion). A single coefficient (H) is sufficient to describe the global statistics, that can be easily estimated by computing statistical moments of order 1 and 2 to obtain the slope of $\zeta(q) = qH$.

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In the *multifractal* case, the scaling function is corrected by a additional quantity $K(q)$, called moment scaling function: $\zeta(q) = qH - K(q)$.

$$\langle \Delta h^q \rangle \sim \Delta x^{\zeta(q)} \quad (1)$$

$$\zeta(q) = qH - K(q) \quad (2)$$

Although three moments of distinct orders are in principle sufficient to establish the curvature of $\zeta(q)$ and assess multiscaling properties, a field with general multifractal symmetries would require to determine all statistical moments (non-integer orders included) to fully describe the statistics. Practically, the estimations of moments of multiple orders can be misleading. Indeed, the estimate of high order moments (typically moment of order > 2) are often biased by multifractal phase transition in the effective $\zeta(q)$: beyond critical values of q , statistical moments may all depend on the largest value in the sample spuriously leading to linear behaviour in the exponent (and diverging as the sample size increases) (Schertzer and Lovejoy, 1993). Therefore, we will cautiously focus on moments of order < 2 that can still provide meaningful information on the shape of $\zeta(q)$. In order to characterize multiscaling, we will use three independent multifractal parameters introduced by Schertzer (1997):

- the *mean fluctuation exponent* $H = \zeta(1)$. Consequence of Eq. (1), the value $H = 0$ is obtained for a field that is strictly scale invariant (conservative), whereas $H \neq 0$ rather corresponds to a fractional integration of a conservative field;
- the *codimension of the mean field* $H - C_1 = \left. \frac{d\zeta(q)}{dq} \right|_{q=1}$ measuring the mean intermittency. If $C_1 = 0$, the field is homogenous. A non-zero value of C_1 indicated intermittency. In the case of topography, C_1 is expected to be close to 0.1 (see the case of Earth, Gagnon et al., 2006);
- the *index of multifractality* $\alpha = \left. \frac{1}{C_1} \frac{d^2\zeta(q)}{dq^2} \right|_{q=1}$ that estimate the curvature of $\zeta(q)$ near $q = 1$. It provides information about the relative variation of intermittency

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around the mean. If $\alpha = 0$, $\zeta(q)$ is linear around $q = 1$ indicating a monofractal behavior.

Consequence of the above definitions, any estimate of H , α and C_1 will rely on computing non-integers order moments around the mean $q = 1$. In the most general multifractal case, this set of three parameters only provides a restricted description of statistics in the neighborhood of the mean (1st order moment) and presumes nothing of the variability that may be revealed by higher order moments. Still there exist a class of multifractal called the *universal multifractals* (Schertzer, 1997) whose associated structure function is fully determined by H , α and C_1 . In other words, for a member of this class, a set of only three parameters is enough to characterize all the statistical moments. Equation (3) indicates the mathematical form expected in the case of *universal multifractality* with a moment scaling function depending only on α and C_1 .

$$K(q) = \frac{C_1}{\alpha - 1} (q^\alpha - q) (0 \leq \alpha \leq 2, 0 \leq C_1). \quad (3)$$

3 Haar fluctuations for irregular signals

Mars is well-known to present huge fluctuations of altitude on its surface despite of its relatively small radius in comparison to earth ($R_{\text{Mars}} = 3390$ km). Indeed Olympus Mons culminates at 22.5 km above the adjacent lowlands. In this context, we have to define fluctuations in order to provide an accurate statistical study. It is well established that data analysis strongly relies on defining the fluctuations at a given scale and location. Of course, the chosen definition must adequately characterize the fluctuations and has to be picked among all the possible definitions depending on the context. As mentioned earlier, the simplest definition is the absolute difference that might not be accurate in some cases. For instance, in the case of fractal field with $H < 0$ (a Gaussian white noise has $H = -1/2$), absolute differences will only characterize fluctuations with high wavenumber instead of the local fluctuations at a given scale. In the case of topography,

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H is expected to be > 0 so absolute differences might be meaningful. Still, it is possible to consider a more general definition of fluctuations that can apply for both cases ($-1 < H < 1$): the Haar fluctuations.

The definition is given by Eq. (A1). For each even numbers $2n$ of along-track consecutive points, we compute the average A of the first n points ($\langle h_i \rangle_1, i = 1 \dots n$) and the average B of the n last points ($\langle h_i \rangle_2, i = n + 1 \dots 2n$). The fluctuation at scale L is defined by the absolute difference $|B - A|$, L being the spherical distance between the first and the last point. On synthetic multifractal series obtained by simulation (Lovejoy and Schertzer, 2012), Haar fluctuations have proven to be strong estimators of the input multifractal parameters. However, Haar fluctuations unlike first differences are more complicated to implement particularly in the case of 1-dimensional irregular series (see Appendix A). For the purpose of this analysis, we computed both 1st difference and Haar fluctuations. Although the two methods produced similar results (as expected since $H > 0$), we noticed a better convergence in the case of Haar fluctuations. Hence, in the result Sect. 4, we will focus on the results obtained by Haar fluctuations.

4 Results

For each scales, we compute moments of several orders $q = 0.1, 0.2 \dots 2$. The scale-dependence of every moment is expected to be a straight line on a log-log plot with a slope predicted by Eq. (2). Figure 1 shows the results (plot on the left). In order to compare the experimental results to simulations, the right plot on Fig. 1 is obtained by performing the same analysis on synthetic series with multifractal parameters $H = 0.5$, $\alpha = 1.7$ and $C_1 = 0.1$. The global aspect of curves on Fig. 1 (right plot) and their resemblance with synthetic curves over a significant range of scales strongly confirm the scaling behavior of topography with a value of H close to 0.5. Still some features need to be discussed. We will focus on scales > 600 m where reasonably straight lines occur (due to the definition of Haar fluctuations, no relevant data can be found outside that range). On Fig. 1 (right plot), statistical moments seems to exhibit two distinct

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scaling regimes. The first regime starts at 10 km to the planet scale 20 000 km covering nearly 4 orders of magnitude. The second one is identified in the range of scale (600 m to 10 km) covering one order of magnitude. Scaling needs to be studied separately for both of these ranges of scale.

4.1 Possible artefacts

Before going any further, we discuss in this section possible artefacts that may cause the observed transition around 10 km:

- Uncertainty due to the accuracy of measurement at small scales: the characteristics presented in Sect. 2.1 indicate that the error on the topography measurement is of a few meters (areoid error + instrument error), lower than the mean fluctuations at the lowest relevant scale (~ 10 m in Fig. 1). Hence, if there is a bias at small scales, it can not be responsible of the transition identified around 10 km, the mean fluctuation at that scale being close to 40 m.
- Due to the quasi-polar orbit of MOLA, the density of measurements depends on latitudes. The minimum is obtained near the equator whereas the maximum is observed near the poles. This could introduced a bias in our results. Indeed, a global analysis might only reflect the statistics of high latitudes regions where the signal is oversampled in comparison to low latitudes. To check that hypothesis, we performed the following: for each contributing fluctuation, we apply a factor $\cos(\theta)$, θ being the latitude of the center of the fluctuations in order to artificially restore the homogeneity. This statistical correction, reduce the weight of the contributions over a range of latitudes to compensate for the oversampling. It turned out that the final result is unaffected by this correction, rejecting the hypothesis of a bias due to an inhomogenous density of measures.
- By construction, all fluctuations have a privileged orientation (Nord-South) due to the orientation of along-track series. We can not exclude that a similar analysis

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would produce a different result with a different along-track direction. This would occur in case of a strong anisotropy of topography at the planetary scale, outside the scope of this analysis.

4.2 Scales < 10 km

5 Over nearly 4 orders of magnitude, beginning from around 10 km up to the planetary scale, the linear behavior of statistical moments up to order 2 is quite accurate, as shown on the right plot from Fig. 1 (area between 4 and 7 on log10 units). The global scaling behavior is established over that range of scales but the multiscaling properties (if they occur) still need to be tested. In order to apply the universal multifractal formalism, we evaluate the slope of each curve in Fig. 1, by computing linear fits. 10 Figure 2 illustrates the results obtained on each curve. As one can see, the linear regression is rather satisfying. For scales > 10 km, each fit provides a scaling exponent that can be studied as a function $\zeta(q)$ (q being the statistical moment order) in order to test the validity of Eqs. (1) and (2). On Fig. 3 (red points), we have plotted the experimental structure function $\zeta(q)$. The apparent curvature of $\zeta(q)$ on Fig. 3 indicates multiscaling. Indeed, in the case of simple scaling behavior, the structure function is expected to be linear: $\zeta(q) = qH$. Instead, the 21 red points on Fig. 3 can not be adjusted by a simple linear fit. 15

The multifractal formalism may be tested by adjusting the experimental structure function $\zeta(q)$ with Eq. (2). On Fig. 3, the red lines shows that the experimental structure is accurately adjusted, indicating that the universal multifractal formalism is well suited to fit the data with only three parameters H , α and C_1 . Moreover, it allows us to estimate the values of the degree of multifractality α and the codimension of the mean C_1 . Those estimates are presented in Table 1. 20

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4.3 Scale > 10 km

Over the range of scales covering only one order of magnitude, the behavior is clearly different. Still topography seems to exhibit scaling behavior over that range. Figure 2 shows the linear fit obtained indicating that the log–log variations of moments according to scale is again satisfyingly linear. However, slopes are significantly steeper: fluctuations decrease faster when the scale decreases, indicating a larger H . The regression slope is computed for each statistical moment in order to study the experimental structure function. Figure 3 (blue points) shows the result. As one can see, the behavior differs from the one analyzed earlier. This time, the q -dependence of the structure function is clearly linear indicating that topography is not multiscaling on that range of scale. Indeed, the estimation of the universal multifractal parameters (see Table 1) leads to a value of C_1 close to 0 (Table 1).

5 Conclusions

Our goal was to validate the accuracy of the universal multifractal formalism to describe the global scaling properties of the martian surface in terms of its topographic intermittency is scaling. From our results, multiscaling seems to occur over a large but restricted range of scale (above 10 km) with. At smaller scale, the topography is still scaling but the symmetry is only monofractal with a parameter $H = 0.75$.

This result is consistent with Aharonson et al. (2001) who obtained similar results by studying the power spectra of two large significant regions of Mars (similar scale break). The power spectrum of a topographic transect is related to the second order moment (square of fourier transform) and is therefore expected to obey a scaling law if the global topography is scaling. Let β be the scaling exponent for the power spectrum. It can be shown that β is related to the structure function for $q = 2$. This relation is given by the following equation:

$$\beta = \zeta(2) + 1 = 2H + 1 - K(2) \quad (4)$$

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Aharonson et al. (2001) found two scaling regimes similar to ours with a transition around 10 km. Our estimates of β for the small and large scale regimes are given in Table 1. Corresponding results from Aharonson et al. (2001) are also reproduced. For scales > 10 km, we found $\beta = 1.85$ to be compared to the estimate of Aharonson et al. (2001) in region 1 (heavily cratered southern terrain) where $\beta = 1.4$ and region 2 (area in the northern lowland) where $\beta = 2.0$. Our global estimate of $\beta = 1.85$ stands between the values obtained by the authors. As we found multiscaling properties that range of scale (> 10 km), we can interpret this difference in terms of intermittency, i.e. local variations of β . Indeed, intermittency is an expected feature of multifractal fields. For scales less than 10 km, our estimate of β differs clearly from the one obtained by Aharonson et al. (2001). Still it might be interesting to notice that the authors found similar value of β in both of these regions, indicating a low intermittency on that range of scale. This is consistent with our previous conclusion about the monofractal nature of scaling on that range of scale.

We added some discussion here on fractured topography (Schmittbuhl et al.)

We demonstrate that a change occurs at 10 km. The generic process can reproduce the statistical simpler monofractal scaling process is occurring a smaller scales. Craterisation is well known to be a fractal process with a single fractal dimension (Rosenburg et al., 2015). We propose that the low scales are dominated by craterisation processes, at the origin of the monofractal scaling law. This process has already been proposed by Aharonson et al. (2001). Future investigation are required to understand the multiplicative cascade processes at large scale.

Appendix A: Haar fluctuations for irregular signals

$$\Delta x(L) = \langle hi \rangle_2 - \langle hi \rangle_1 \quad (\text{A1})$$

A1 Irregularity

The above definition for haar fluctuations implicitly requires that along-track points are regularly spaced. However, the actual MOLA datasets exhibits different kind of irregularities, mainly due the occurrence of clouds of various length and random dysfunctionnement of the instrument that might introduce a bias in our analysis. Figure 4 provides typical examples of such irregularities in the MOLA along track series and Fig. 5 shows the distribution of distance between two along-track points computed among the entire MOLA dataset. It clearly indicates that most of the points are ~ 300 m spaced, as expected but a subsequent numbers of irregularities occur (1 point missing, 2 points missing, etc.).

In order to take this into account, we define a adjustable quality criterium on the Haar fluctuations (see Lovejoy, 2014): for a given fluctuation at scale L we compute the spherical length of first half L_1 and the spherical length of the second half L_2 . In the case of perfect regularity L_1 is exactly equal de L_2 . If the measures are non-equally spaced, L_1 might be inferior or superior to L_2 . The degree of irregularity may be estimated by a ratio R defined by Eq. (A2), equal to 0 if the measures are regular and close to $\frac{1}{2}$ in case of extreme irregularity.

$$R = \left| \frac{1}{2} - \frac{L_1}{L_1 + L_2} \right| \quad (\text{A2})$$

The next step is to define a threshold to be applied on the ratio so that irregular fluctuations can be excluded of the global analysis. The threshold has to to be chosen carefully: a very restrictive threshold might be damageable by excluding a lot of fluctuations. Indeed a huge amount of fluctuations is needed for a statistical purpose. On the other hand, an un-restrictive threshold might involve unrelevant fluctuations in the calculation. That also may introduce a bias.

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A2 Choice of a threshold

In this paragraph, we try to quantify how the choice of a threshold might affect the analysis. For that purpose, we have analyzed multifractal simulations after having manually introduced holes of different kind in the data. To produce relevant statistics, 1000 series of universal multifractal series have been generated using the simulation technique called the Fractionally Integrated Flux (FIF) developed by Schertzer and Lovejoy (1987) (<http://www.physics.mcgill.ca/>). Each series is composed of 100 000 equally-spaced points. The universal scaling parameters used for this simulation are $H = 0.5$, $\alpha = 1.7$ and $C_1 = 0.1$. That particular choice relies on the assumption that scaling laws on Mars might be similar to the ones observed on Earth (see Gagnon et al., 2006). First step is to introduce irregularities in order to reproduce the MOLA distribution of distance between consecutive points plotted on Fig. 5. This is obtained by suppressing data from the synthetic series. The accurate number of artificial holes of different kind (1pt missing, 2 pts missing, etc.) is adjusted by comparison with the actual distribution of holes plotted on Fig. 5. This figure compares the effective irregularity of MOLA (blue line) and the synthetic irregularity distribution manually introduced on simulations (red points). The kind and distribution of irregularities is realistic if red points coincide with maxima of the blue lines. As one can see on Fig. 5, the correspondence between simulation and actual data is satisfying.

We can now compute Haar fluctuations on both biased and unbiased data and see how the choice of a threshold impacts the analysis on the biased series. We evaluate statistical moment of order $q = 0.5$, $q = 1$, $q = 1.5$ and $q = 2$ at different scales with different values of threshold R . Figure 6 shows how the biased moment deviate from the unbiased ones depending on the threshold. As one can see, the difference between blue dots and red dots on Fig. 6 is minim, regardless of the threshold. Our conclusion is that in the case of a huge dataset and given a relatively small density of missing points, the statistical multifractal properties are not significantly affected by high-ratios fluctuations. Hence, we are left to chose the least restrictive ratio, in order to involve

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as many fluctuations as possible in the calculation and optimize the convergence. In the following analysis, the ratio is fixed at $R = 0.5$ (completely unrestrictive). The total numbers of fluctuations computed to perform this analysis is over a billions that spread into 74 bins of scales, equally spaced logarithmically from the largest scale possible (10 000 km, the half- circumference of the planet) down to the meter scale. Figure 7 shows the distribution of Haar fluctuations according to scale. Under 600 m ($\sim 10^{2.7}$), the number of fluctuations decrease really fast, consequence of the definition of Haar fluctuations.

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References

- Aharonson, O., Zuber, M. T., and Rothman, D. H.: Statistics of Mars’ topography from the Mars orbiter laser altimeter: slopes, correlations, and physical Models, *J. Geophys. Res.-Planet.*, 106, 23723–23735, doi:10.1029/2000JE001403, 2001. 2, 3, 11, 12, 18
- Baldassarri, A., Montuori, M., Prieto-Ballesteros, O., and Manrubia, S. C.: Fractal properties of isolines at varying altitude revealing different dominant geological processes on Earth, *J. Geophys. Res.*, 113, E09002, doi:10.1029/2007JE003066, 2008. 3
- Gagnon, J.-S., Lovejoy, S., and Schertzer, D.: Multifractal earth topography, *Nonlin. Processes Geophys.*, 13, 541–570, doi:10.5194/npg-13-541-2006, 2006. 3, 4, 6, 14
- Gilbert, L.: Are Topographic Data Sets Fractal?, in: *Fractals in Geophysics*, edited by: Scholz, C. H. and Mandelbrot, B. B., *Pure and Applied Geophysics* 241–254, doi:10.1007/978-3-0348-6389-6_13, 1989. 3
- Kreslavsky, M. A. and Head, J. W.: Kilometer-scale roughness of Mars: results from MOLA data analysis, *J. Geophys. Res.-Planet.*, 105, 26695–26711, doi:10.1029/2000JE001259, 2000. 2
- Lavallee, D., Lovejoy, S., Schertzer, D., and Ladoy, P.: Nonlinear variability and landscape topography: analysis and simulation, in: *Fractals in Geography*, edited by: De Cola, L. and Lam, N., PTR, Prentice Hall, Englewood, N.J., 158–192, 1993. 3

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- Lemoine, F. G., Smith, D. E., Rowlands, D. D., Zuber, M. T., Neumann, G. A., Chinn, D. S., and Pavlis, D. E.: An improved solution of the gravity field of Mars (GMM-2B) from Mars Global Surveyor, *J. Geophys. Res.-Planet.*, 106, 23359–23376, doi:10.1029/2000JE001426, 2001. 4
- 5 Lovejoy, S.: A voyage through scales, a missing quadrillion and why the climate is not what you expect, *Clim. Dynam.*, 44, 3187–3210, doi:10.1007/s00382-014-2324-0, 2014. 13
- Lovejoy, S. and Schertzer, D.: Haar wavelets, fluctuations and structure functions: convenient choices for geophysics, *Nonlin. Processes Geophys.*, 19, 513–527, doi:10.5194/npg-19-513-2012, 2012. 8
- 10 Lovejoy, S. and Schertzer, D.: *The Weather and Climate: Emergent Laws and Multifractal Cascades*, Cambridge, University Press, 496 pp., 2013 5
- Malamud, B. D. and Turcotte, D. L.: Wavelet analyses of Mars polar topography, *J. Geophys. Res.-Planet.*, 106, 17497–17504, doi:10.1029/2000JE001333, 2001. 3
- Mandelbrot, B.: How long is the coast of Britain? Statistical self-similarity and fractional dimension, *Science*, 156, 636–638, doi:10.1126/science.156.3775.636, 1967. 3
- 15 Nikora, V. and Goring, D.: Mars topography: bulk statistics and spectral scaling, *Chaos Soliton. Fract.*, 19, 427–439, doi:10.1016/S0960-0779(03)00054-7, 2004. 3
- Orosei, R., Bianchi, R., Coradini, A., Espinasse, S., Federico, C., Ferriccioni, A., and Gavrishin, A. I.: Self-affine behavior of Martian topography at kilometer scale from Mars orbiter laser altimeter data, *J. Geophys. Res.*, 108, 8023, doi:10.1029/2002JE001883, 2003. 3
- 20 Pommerol, A., Chakraborty, S., and Thomas, N.: Comparative study of the surface roughness of the Moon, Mars and Mercury, *Planet. Space Sci.*, 73, 287–293, 2012. 3
- Rosenburg, M. A., Aharonson, O., Head, J. W., Kreslavsky, M. A., Mazarico, E., Neumann, G. A., Smith, D. E., Torrence, M. H., and Zuber, M. T.: Global surface slopes and roughness of the Moon from the Lunar Orbiter Laser Altimeter, *J. Geophys. Res.*, 116, E02001, doi:10.1029/2010JE003716, 2011. 3
- 25 Rosenburg, M. A., Aharonson, O., and Sari, R.: Topographic power spectra of cratered terrains: theory and application to the Moon, *J. Geophys. Res.-Planet.*, 120, 177–194, doi:10.1002/2014JE004746, 2015. 3, 12
- 30 Schertzer, D. and Lovejoy, S.: Physical modeling and analysis of rain and clouds by anisotropic scaling multiplicative processes, *J. Geophys. Res.-Atmos.*, 92, 9693–9714, doi:10.1029/JD092iD08p09693, 1987. 3, 14

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Schertzer, D. and Lovejoy, S.: Generic multifractal phase transitions and self-organized criticality, *Cellular Automata: prospects in astronomy and astrophysics*, edited by: Perchang, J. M., Lejeune, A., World Scientific, Hong Kong, 216–227, 1993. 6

Schertzer, D. and Lovejoy, S.: Universal multifractals do exist!: Comments on A statistical analysis of mesoscale rainfall as a random cascade, *J. Appl. Meteorol.*, 36, 1296–1303, 1997. 6, 7

Shepard, M. K., Campbell, B. A., Bulmer, M. H., Farr, T. G., Gaddis, L. R., and Plaut, J. J.: The roughness of natural terrain: a planetary and remote sensing perspective, *J. Geophys. Res.-Planet.*, 106, 32777–32795, doi:10.1029/2000JE001429, 2001. 2, 3

Smith, D. E., Zuber, M. T., Frey, H. V., Garvin, J. B., Head, J. W., Muhleman, D. O., Pettengill, G. H., Phillips, R. J., Solomon, S. C., Zwally, H. J., Banerdt, W. B., Duxbury, T. C., Golombek, M. P., Lemoine, F. G., Neumann, G. A., Rowlands, D. D., Aharonson, O., Ford, P. G., Ivanov, A. B., Johnson, C. L., McGovern, P. J., Abshire, J. B., Afzal, R. S., and Sun, X.: Mars orbiter laser altimeter: experiment summary after the first year of global mapping of Mars, *J. Geophys. Res.-Planet.*, 106, 23689–23722, doi:10.1029/2000JE001364, 2001. 4

Turcotte, D. L.: *Fractals and chaos in geology and geophysics*, Cambridge University Press, Cambridge, England, 1997. 3

Vidal Vázquez, E., García Moreno, R., Miranda, J. G. V., Díaz, M. C., Saá Requejo, A., Paz Ferreiro, J., and Tarquis, A. M.: Assessing soil surface roughness decay during simulated rainfall by multifractal analysis, *Nonlin. Processes Geophys.*, 15, 457–468, doi:10.5194/npg-15-457-2008, 2008. 3

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Table 1. Estimates of the fractal and multifractal parameters for scales below and above 10 km. The last two columns were taken from Aharonson et al. (2001). Region 1 is a heavily cratered southern terrain. Region 2 is an area in the northern lowland.

scale	H	α	C_1	$A(2)$	β	$\beta(\text{reg.1})$	$\beta(\text{reg.2})$
> 10 km	0.52	1.86	0.11	0.85	1.85	1.4	2.0
< 10 km	0.76	-0.16	-0.02	1.43	2.43	3.4	3.4

the irrelevant values for
alpha and C1 have been
removed

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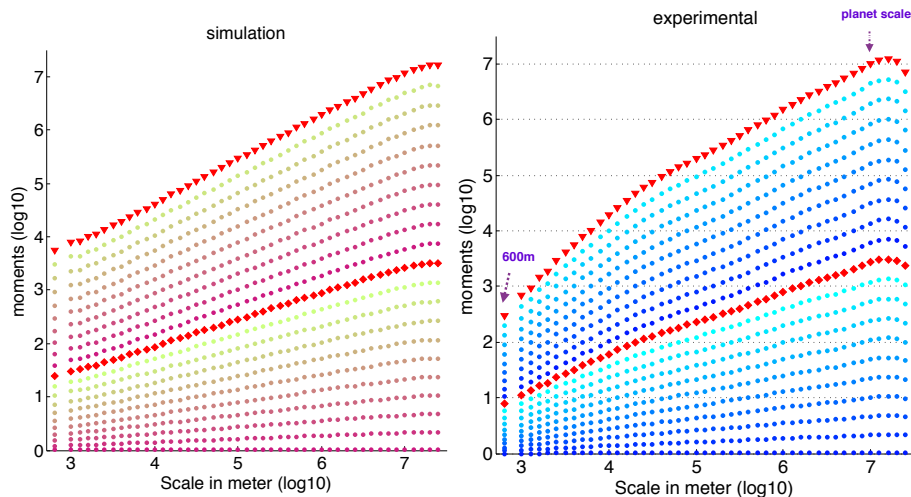


Figure 1. Statistical moments of several orders (from 0.1 to 2) as a function of scale for simulated series (left plot) and for actual MOLA data (right plot). In both plot, red diamonds (respectively red triangles) stand for moment of order 1 (respectively order 2). Integer ordered moments are also computed in order to test the multifractal formalism. They are plotted in shades of red (left plot) and blue (right plot).

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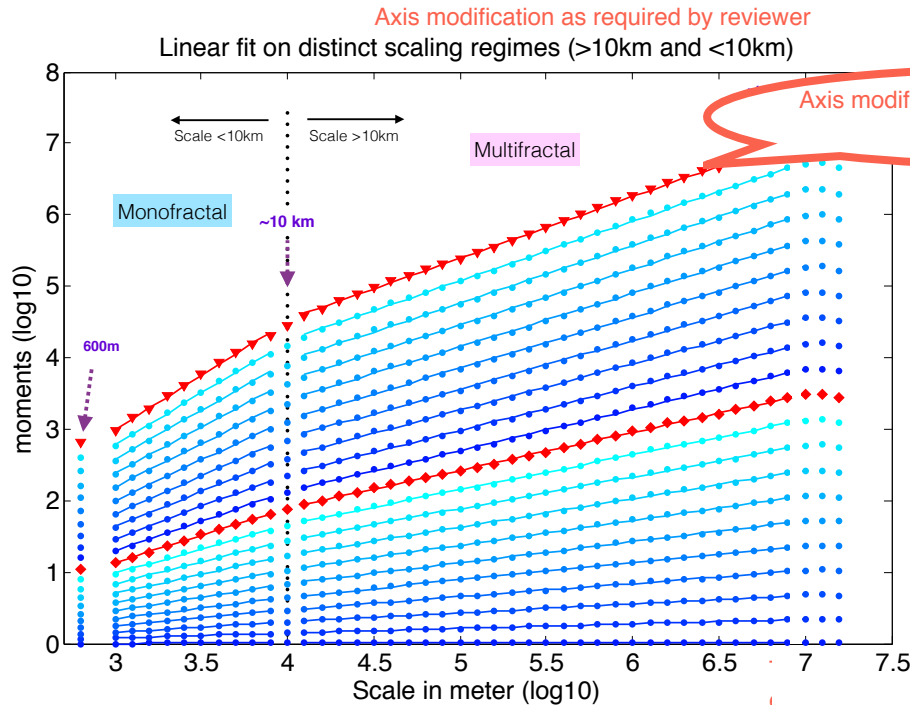


Figure 2. Linear fit on the two different scaling regimes (inferior and superior to 10 km) for every 21 statistical moments from 0.1 to 2. The color code is the same as in Fig. 1. Quantities on that plot are scaling according to the moment scaling function $\xi(q)$.

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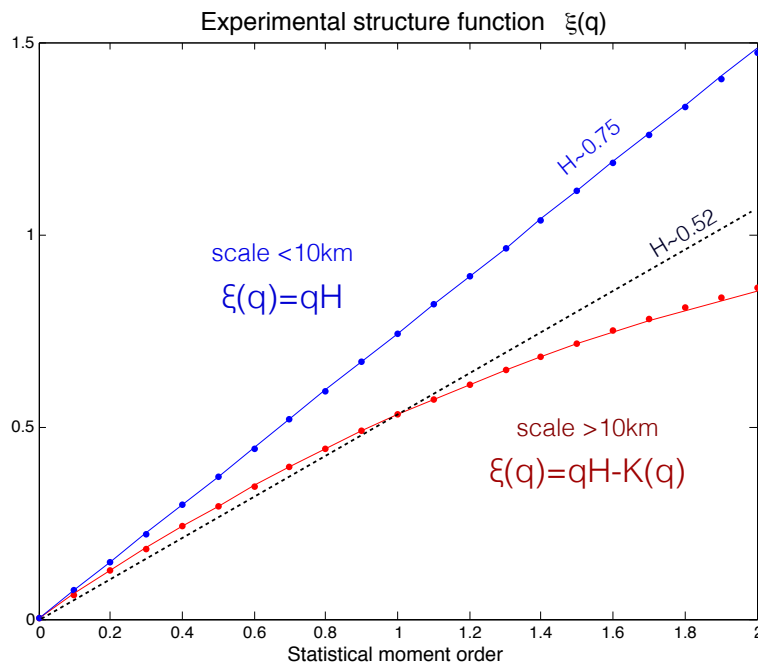


Figure 3. Theoretical structure function $\zeta(q)$ combining the 21 linear fits shown over Fig. 2. Red points (resp. blue points) correspond to the range of scales larger (resp. smaller) than 10 km.

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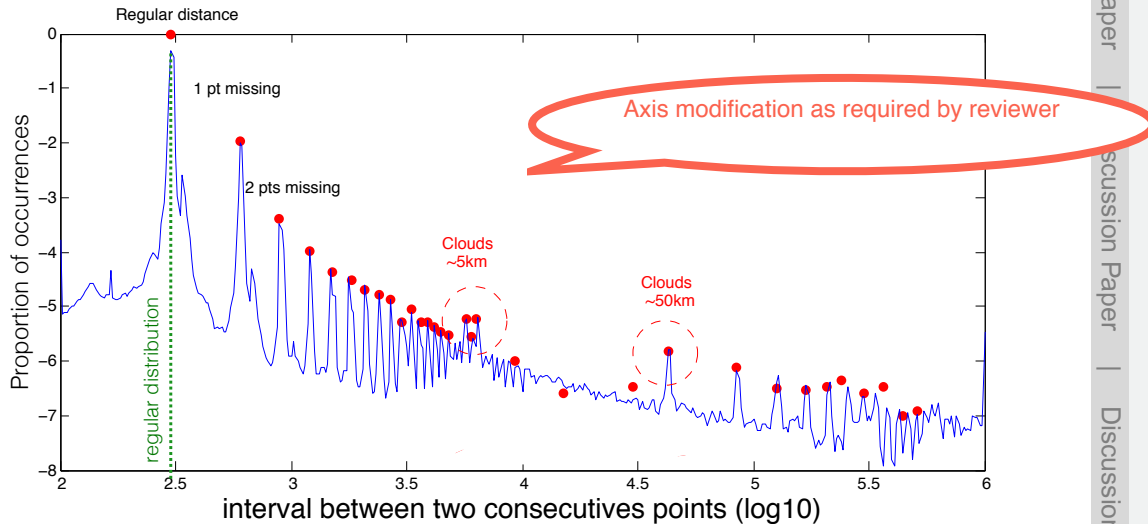


Figure 5. Comparison between effective irregularities in the MOLA database and artificial irregularities on simulations. Irregularities are expressed as a proportion of occurrences from the maximum possible in log10 space (0 means all fluctuations are observed). The blue line represents the effective irregularity of actual data. Red points are obtained from the synthetic irregularities on multifractal simulations. The good agreement between maxima of the blue curve and red points indicates that effective irregularity is accurately reproduced by simulations.

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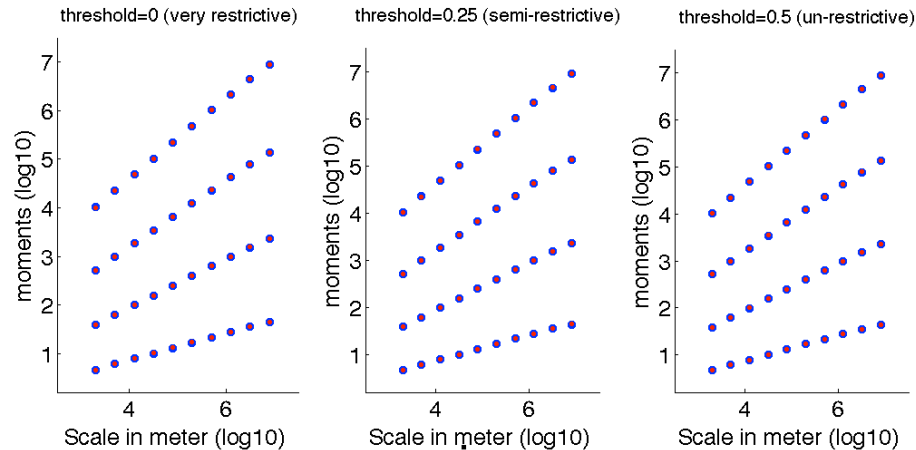


Figure 6. Statistical moments of order $q = 0.5$, $q = 1$, $q = 1.5$ and $q = 2$ computed for biased and unbiased synthetic data at different scales. From left to right, the threshold is $R = 0, 0.25, 0.5$ from the more to the less restrictive. Blue points stand for the regular reference data. Red points represent irregular series (all ratios). Red points are satisfyingly centered inside blue points, regardless of the threshold. Therefore, the kind of irregularity that occurs on real data does not significantly impact the quality of Haar analysis.

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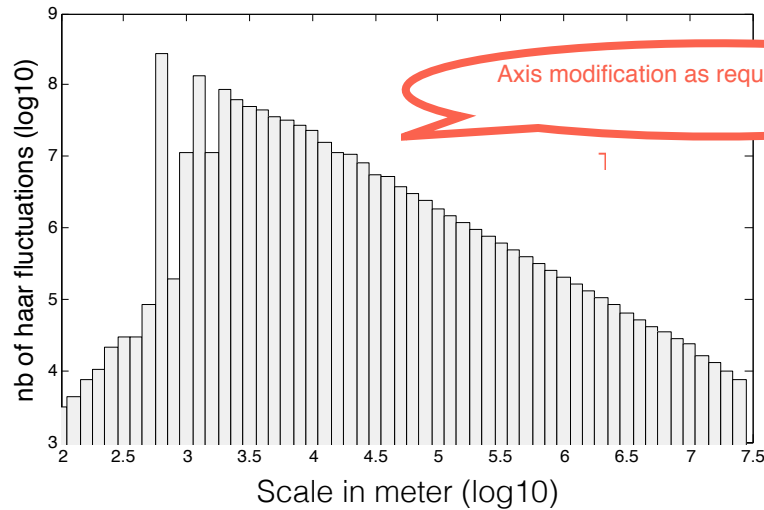


Figure 7. Distributions of Haar fluctuations depending on the scale.

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